

graviton

$$J = 2$$

graviton MASS

It is likely that the graviton is massless. More than fifty years ago Van Dam and Veltman (VANDAM 70), Iwasaki (IWASAKI 70), and Zakharov (ZAKHAROV 70) almost simultaneously showed that in the linear approximation a theory with a finite graviton mass does not approach GR as the mass approaches zero. Attempts have been made to evade this "vDVZ discontinuity" by invoking modified gravity or nonlinear theory by De Rahm (DE-RHAM 17) and others. More recently, the analysis of gravitational wave dispersion has led to bounds that are largely independent of the underlying model, even if not the strongest. We quote the best of these as our best limit.

Experimental limits have been set based on a Yukawa potential (YUKA), dispersion relation (DISP), or other modified gravity theories (MGRV).

The following conversions are useful: $1 \text{ eV} = 1.783 \times 10^{-33} \text{ g} = 1.957 \times 10^{-6} m_e$; $\lambda_C = (1.973 \times 10^{-7} \text{ m}) \times (1 \text{ eV}/m_g)$.

VALUE (eV)	DOCUMENT ID	TECN	COMMENT
<5 × 10⁻²³	1 ABBOTT	19 DISP	LIGO Virgo catalog GWTC-1
● ● ● We do not use the following data for averages, fits, limits, etc. ● ● ●			
<3.2 × 10 ⁻²³	2 BERNUS	20 YUKA	Planetary ephemeris INPOP19a
<2 × 10 ⁻²⁸	3 SHAO	20 DISP	Binary pulsar Galileon radiation
<7 × 10 ⁻²³	4 BERNUS	19 YUKA	Planetary ephemeris INPOP17b
<3.1 × 10 ⁻²⁰	5 MIAO	19 DISP	Binary pulsar orbital decay rate
<1.4 × 10 ⁻²⁹	6 DESAI	18 YUKA	Gal cluster Abell 1689
<5 × 10 ⁻³⁰	7 GUPTA	18 YUKA	Using SPT-SZ
<3 × 10 ⁻³⁰	7 GUPTA	18 YUKA	Using Planck all-sky SZ
<1.3 × 10 ⁻²⁹	7 GUPTA	18 YUKA	Using redMaPPer SDSS-DR8
<6 × 10 ⁻³⁰	8 RANA	18 YUKA	Weak lensing in massive clusters
<8 × 10 ⁻³⁰	9 RANA	18 YUKA	SZ effect in massive clusters
<1.0 × 10 ⁻²³	10 WILL	18 YUKA	Perihelion advances of planets
<7 × 10 ⁻²³	1 ABBOTT	17 DISP	Combined dispersion limit from three BH mergers
<1.2 × 10 ⁻²²	1 ABBOTT	16 DISP	Combined dispersion limit from two BH mergers
<2.9 × 10 ⁻²¹	11 ZAKHAROV	16 YUKA	S2 star orbit
<5 × 10 ⁻²³	12 BRITO	13 MGRV	Spinning black holes bounds
<6 × 10 ⁻³²	13 GRUZINOV	05 MGRV	Solar System observations
<6 × 10 ⁻³²	14 CHOUDHURY	04 YUKA	Weak gravitational lensing
<9.0 × 10 ⁻³⁴	15 GERSHTEIN	04 MGRV	From Ω_{tot} value assuming RTG
<8 × 10 ⁻²⁰	16,17 FINN	02 DISP	Binary pulsar orbital period decrease
<7 × 10 ⁻²³	TALMADGE	88 YUKA	Solar system planetary astrometric data
<1.3 × 10 ⁻²⁹	18 GOLDHABER	74 YUKA	Rich clusters
<7 × 10 ⁻²⁸	HARE	73 YUKA	Galaxy
<8 × 10 ⁴	HARE	73 YUKA	2 γ decay

- ¹ ABBOTT 19, ABBOTT 17, and ABBOTT 16 limits assume a dispersion relation for gravitational waves modified relative to GR.
- ² BERNUS 20 use the latest solution of the ephemeris INPOP (19a) in order to improve the constraint in BERNUS 19 on the existence of a Yukawa suppression to the Newtonian potential, generically associated to a gravitons mass.
- ³ SHAO 20 sets limit, 95% CL, based on non-observation of excess gravitational radiation in 14 well-timed binary pulsars in the context of the cubic Galileon model.
- ⁴ BERNUS 19 use the planetary ephemeris INPOP 17b to constraint the existence of a Yukawa suppression to the Newtonian potential, generically associated to a gravitons mass.
- ⁵ MIAO 19 90% CL limit is based on orbital period decay rates of 9 binary pulsars using a Bayesian prior uniform in graviton mass. Limit becomes $< 5.2 \times 10^{-21}$ eV for a prior uniform in $\ln(m_g)$.
- ⁶ DESAI 18 limit based on dynamical mass models of galaxy cluster Abell 1689.
- ⁷ GUPTA 18 obtains graviton mass limits using stacked clusters from 3 disparate surveys.
- ⁸ RANA 18 limit, 68% CL, obtained using weak lensing mass profiles out to the radius at which the cluster density falls to 200 times the critical density of the Universe. Limit is based on the fractional change between Newtonian and Yukawa accelerations for the 50 most massive galaxy clusters in the Local Cluster Substructure Survey. Limits for other CL's and other density cuts are also given.
- ⁹ RANA 18 limit, 68% CL, obtained using mass measurements via the SZ effect out to the radius at which the cluster density falls to 500 times the critical density of the Universe for 182 optically confirmed galaxy clusters in an Altacama Cosmology Telescope survey. Limits for other CL's and other density cuts are also given.
- ¹⁰ WILL 18 limit from perihelion advances of the planets, notably Earth, Mars, and Saturn. Alternate analysis yields $< 6 \times 10^{-24}$.
- ¹¹ ZAKHAROV 16 constrains range of Yukawa gravity interaction from S2 star orbit about black hole at Galactic center. The limit is $< 2.9 \times 10^{-21}$ eV for $\delta = 100$.
- ¹² BRITO 13 explore massive graviton (spin-2) fluctuations around rotating black holes.
- ¹³ GRUZINOV 05 uses the DGP model (DVALI 00) showing that non-perturbative effects restore continuity with Einstein's equations as the graviton mass approaches zero, then bases his limit on Solar System observations.
- ¹⁴ CHOUDHURY 04 concludes from a study of weak-lensing data that masses heavier than about the inverse of 100 Mpc seem to be ruled out if the gravitation field has the Yukawa form.
- ¹⁵ GERSHTEIN 04 use non-Einstein field relativistic theory of gravity (RTG), with a massive graviton, to obtain the 95% CL mass limit implied by the value of $\Omega_{tot} = 1.02 \pm 0.02$ current at the time of publication.
- ¹⁶ FINN 02 analyze the orbital decay rates of PSR B1913+16 and PSR B1534+12 with a possible graviton mass as a parameter. The combined frequentist mass limit is at 90%CL.
- ¹⁷ As of 2020, limits on dP/dt are now about 0.1% (see T. Damour, "Experimental tests of gravitational theory," in this *Review*).
- ¹⁸ GOLDHABER 74 establish this limit considering the binding of galactic clusters, corrected to Planck $h_0 = 0.67$.

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